SOAR Optical Imager (SOI)

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The SOAR Optical Imager (SOI) is a bent-Cassegrain mounted optical imager using a mini-mosaic of two E2V 2k x 4k CCDs to cover a 5.26 arcminute square field of view at a scale of 0.077"/pixel. It was designed, built, and integrated at SOAR by an NOAO/CTIO team lead by Drs. Alistair Walker and Hugo Schwarz.

Two views of the SOI mounted at one of the bent Cassegrain foci of the SOAR telescope. The smaller blue cylinder is the CCD mosaic dewar and is flanked by the two rectangular Leach controllers, the round black part is the cable wrap, and the large blue tube bolted onto the telescope is the main body of the instrument holding the linear ADC prisms, one of which can move over a distance of nearly one meter. The drawing shows the basic lay-out of the instrument in a cut-out view. From left to right the light encounters: ADC prisms, optics module, two filter slides, the shutter, and finally the CCD mosaic.

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SOAR Optical Imager Manual

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SOI Overview

At least read this!

Figure 1: Two views of the SOI mounted at one of the bent Cassegrain foci of the SOAR telescope. The smaller blue cylinder is the CCD mosaic dewar and is flanked by the two rectangular Leach controllers, the round black part is the cable wrap, and the large blue tube bolted onto the telescope is the main body of the instrument holding the linear ADC prisms, one of which can move over a distance of nearly one meter. The drawing shows the basic lay-out of the instrument in a cut-out view. From left to right the light encounters: ADC prisms, optics module, two filter slides, the shutter, and finally the CCD mosaic.


Optics:
The SOI optics consist of a 6 element focal reducer and field corrector, that converts the f/16.63 beam of
the SOAR Telescope to f/9.82, preceded by a linear, "trombone style", Atmospheric Dispersion Corrector (ADC). These optics were designed to deliver images of < 0.18 arcsec FWHM (equivalent to 80% encircled energy within D80 < 0.27 arcsec) at the zenith and <0.34 arcsec FWHM (D80 < 0.51 arcsec) at 70deg zenith distance, in each of the U, B, V, R, I broad band filters, and over the entire field. The glasses selected, and the use of SOLGEL over MgF2 coatings on all external surfaces ensure high transmission over the entire 310nm-1050nm passband. The measured transmission curve is shown below in Figure 2. Note that the transmission below 400nm is an underestimate of the true transmission due to limitations of the measuring device, and should be considered a lower limit.

![SOI Transmission](image)

Figure 2. SOI measured optical transmission. See note in text.

**Filters:**
Filters are mounted in two filter cartridges each with space for 4 filters, plus one blank position, so that up to 8 filters can be installed in the instrument at any time. Although the cartridges were designed to be easily removable, the process of changing them out still requires at least 30 minutes, and we currently have no spare cartridges. Thus it is strongly recommended that users design their programs to avoid filter changes during the night. Filters may be up to 100mm square and up to 10mm thick. Smaller filters may be accommodated with special adapters but must be at least 64mm square to avoid vignetting. To see the list of the currently available filters, together with transmission curves, click [HERE](#). Filters from the [CTIO filter list](#) are also available.

**CCDs:**
The SOI focal plane is imaged onto a mini-mosaic of two CCD's with the following properties:
CCD mounting gap:
The two CCDs in the SOI are mounted with their long sides parallel and spaced 102 pixels apart, resulting in a 7.8" gap between the individual CCD images. This image of 30 Doradus shows how a single SOI image looks. The gap can be filled by taking dithered images. We recommend taking at least 3 images with 10" steps to produce a complete combined image with no gap.

Flat Fields and Fringing:
The CCDs show fringing in the $I$, $i'$ and $z'$ bands (see example here) [42] at the 10%, 3% and 12% level, respectively. In the $U$ band some hatched structure is seen but this flat fields out perfectly. Here are examples of flatfields at the Bessell $U$ [43], $B$ [44], $V$ [45], $R$ [46], and $I$ [47] bands (click the band to see the link).

Ghosts, stray light etc...:
The use of interference filters - which select their passband by reflecting the rest of the light - results in haloes around bright sources. These haloes look like images of the entrance pupil of the telescope with the secondary mirror spider clearly visible, especially when the seeing is good. We have some stray light from encoders in the SOI rotator on the CCDs in the $I$ band. The level is about 0.2 ADU per second, or less than 1% of the dark sky background. We are working on eliminating this stray light. A complete report on these and other optical effects in the SOI is given in a report, available here.
**Read-out:**
The chips are read out by two amplifiers each. Various read-out and binning modes will be available. Note that a seeing disk of 0.45" is still well-sampled by 2x2 binned pixels, so that only for the very best seeing and using the tip-tilt mode should the CCDs be used in unbinned mode. At present, the tip-tilt mode is not yet offered. SOI is no longer offered in slow read-out mode (as of 2008 Nov).

**Focusing:**
Focusing of the SOI is done by charge shifting on the CCD, and moving the telescope focus between steps by automatic software. The number of steps, step size, and exposure time per subexposure can be chosen by the observer.

**Guider:**
The instrument incorporates a fast readout CCD guide camera which can be used to drive the telescope's tertiary mirror in order to partially compensate atmospheric Tip and Tilt, and correct telescope tracking jitter. The actual performance of this system has yet to be determined, but it is expected that Tip-Tilt correction will be possible with guide stars as faint as R=14, with useful guiding performance down to R=19. The field of view of the guide camera is only 7" square, however, the probe can be placed anywhere within a 12 arcmin diameter patrol field. Normally the probe should be kept outside the central 6x6 arcmin to avoid blocking part of the science field. However, when the best possible Tip-Tilt correction is required it may be desireable to select a guide star close to the science target accepting the obstruction of parts of the field.

**Introduction to SOI**

**Instrument Overview**
The SOAR Optical Imager is a mini-mosaic comprised of two 2048x4096 E2V CCDs (4096x4096 pixels total). At the bent-Cassegrain port, the pixels subtend 0.0767 arcsec on the sky, providing a field of view 5.25 arcmin on a side.

The CCDs are read by two Leach controllers through 4 amplifiers (2 amps per CCD). Depending on binning and the gain setting, the CCDs can be read in as little as 6.1 seconds (4x4 fast readout) to as long as 106.0 seconds (1x1 slow readout). Please see the table given in the SOI Overview page for a more detailed description. The data are taken and examined via vncviewers on the SOI data acquisition computer (currently known as soaric1). From soaric1, one can transfer the data to their home institution.

Unbinned SOI images plus overscan and header information are approximately 37 Mbytes each. Most normal observations are performed in 2x2 binning mode with a binned pixel size of ~0.15 arcsec and are approximately 9 Mbytes each. A typical night produces about 2 Gbytes of data and easily transferred over the internet. This is the preferred method of the SOAR partners. If this is unfeasible, please contact Sean Points prior to your run so that other options can be discussed.

The SOI filter cartridges contain space for 4 filters, plus one blank position, so that up to 8 filters can be installed in the instrument at the same time. Filters may be up to 100mm square and up to 10mm thick. Smaller filters may be accommodated with special adapters but must be at least 64mm square to avoid vignetting.
Philosophy and Structure of this Manual

This manual is intended for an observer planning to use the SOAR Optical Imager (SOI). It is not intended to serve as a hardware or software reference document describing the inner working of SOI, although some details at that level may appear to help the observer plan observing strategies. Also, we assume that the observer is already familiar with CCD cameras, observations, and data reductions.

The SOI Overview [12] is at the front of this manual. If you've read this far, and don't plan to read any further, be sure you understand the SOI Overview [12] pages.

This manual follows the outlines of the KPNO and CTIO Mosaic imager manuals. Therefore, we have also adopted the philosophy of bald-faced plagarism and have kept structure of those documents. We have made changes to reflect the differences of the SOI mini-mosaic to the KPNO and CTIO Mosaic imagers. We will update this version as improved data and information become available.

Development of the SOAR Optical Imager system is a continuing process. Throughout the lifetime of the instrument, filters will be added, old ones replaced, and software enhanced. This manual represents the status as of the date on the cover page. We expect to revise the manual occasionally to include information gained during engineering runs, as well as to reflect new filters.

Supplemental Information

Other useful information regarding the use of Mosaic, CCDs, and observing and reduction software can be found at:

- Direct Imaging Manual [48]
- KPNO Mosaic Web Pages [49]

The SOI Hardware

The SOI CCDs

The SOAR Optical Imager (SOI) features two 2048 (serial or pixels/row) x 4096 (parallel or pixels/column) 15 micron per pixel CCDs arranged as a 4096 x 4096 pixel detector. The CCDs are read out through two amplifiers per chip simultaneously using an SDSU-2 Leach controller. The gap between the CCDs is kept to about 102 pixels in the row direction (see Figure 2 showing an image labeled with chip numbers). We have populated the SOI with thinned, back illuminated E2V CCDs. These chips have only minor flaws which have little effect on their scientific performance.
Figure 2: A flat-field (R band) map of the SOI CCDs with their FITS extensions (im1, im2, im3, and im4) labeled accordingly.

Figure 3 shows the average QE of the SOI CCDs.
The Dewar

The SOI dewar contains a highly polished aluminum vessel (11.6 liters) that can be filled with LN2 to 50% capacity via an axial tube and a hold time of 30+ hours. The LN2 can is connected to the CCD mount by a copper braid which has had its thermal resistance optimized. The CCD mount is a kinematic design, with a frame coupled to the dewar front face by G-10 stand-offs. The frame in turn carries the CCD mount and the heater, which is servoed to keep the CCD at a constant temperature to within 0.2 degrees C. The SOI dewar is the smaller blue cylinder on the right-hand side of Figure 4. It is filled by an observing technician at the start of each night.

Figure 3: The average E2V CCD QEs.
The Leach Controllers

The two CCDs are read out by an SDSU-3 Leach controller (see Figure 4), connected to a PC running GNU/Linux through a PCI board in a PCI-GNU/Linux-LabVIEW environment. The read-out time using two amplifiers per CCD is 20 seconds in unbinned mode with a 4.4e read noise. Other modes of operation are given in the SOI Overview document at the beginning of this manual. Data values are stored as 16-bit unsigned integers.

The CCD Shutter

The shutter is a single blade "focal plane" type driven by a dc servomotor via a cogged belt. It is low profile (4.7mm) and high performance. Minimum exposure time is less than 200 msec and repeatability below 0.5 msec. The drive profile is trapezoidal. The shutter has been subject to extensive lifetime tests without degradation.

The Filter Wheels

Filters are mounted in two filter cartrides each with space for 4 filters, plus one blank position, so that up to 8 filters can be installed in the instrument at any time. Although the cartrides were designed to be easily removeable, the process of changing them out still requires at least 30 minutes, and we currently have no spare cartrides. Thus it is strongly recommended that users design their programs to avoid filter changes during the night. Filters may be up to 100mm square and up to 10mm thick. Smaller filters may be accommodated with special adapters but must be at least 64mm square to avoid vignetting. To see the list of the currently available filters, together with transmission curves, click HERE [6]. Filters from the CTIO filter list [50] are also available.

The Tip-tilt Guider

The tip-tilt guider consists of a probe containing a miniature (10 mm side) right angled prism which
diverts a small portion of the input beam to a second prism, where the beam relay optics send it to the tip-tilt guider detector. The prism size and probe arm cross-section are small (the latter has a cross-section of only 5 mm) because the prism and probe arm may need to occult the science beam if the guide star is desired, or can only be found, close to the field center. The probe assembly is mounted on a Parker-Daedalus XY-stage, with 0.5 micron resolution. It is also possible to fine-tune the focus independently from the telescope focus. The detector system is a thinned, quad-readout Marconi CCD-39 that is thermoelectrically cooled and operated by an SDSU Leach controller in frame-transfer mode.

The Rotator and Cable-wrap

The rotator mechanism is driven by a servo motor, directly coupled to a Harmonic Drive gear reduction. This turns a bronze friction wheel against a larger steel wheel which houses a 4-point ball bearing that is attached to the rotating part of the instrument. Position sensing is achieved by a tape encoder located inside the bearing housing. It has a resolution of 0.16 arcsec, more than needed. The rotator has a maximum rotation speed of 0.48 degrees per second and has redundant limit switches for safety. The cable-wrap consists of a IGUS energy chain system, specially manufactured for this application. It is lubricated with Teflon to reduce friction.

The Focal Reducer Optics

All elements of the focal reducer are held in individual cells. The triplet is held by 8 preloaded spring elements with hard points only in the axial direction. The compression of the spring elements was calculated to achieve an acceptable centering. The other two cells are of the "finger" type commonly used in CTIO instruments.

The Atmospheric Dispersion Corrector

The atmospheric dispersion corrector (ADC) incorporated into the SOI is a two-prism linear "trombone-like" design, such as that adopted for the FORS instrument on the VLT. This design is particularly well suited for and alt-azimuth telescope as the angle between the image plane (i.e., the telescope altitude plane) and the atmospheric dispersion axis remains constant for all zenith distances, thus avoiding the need for prism rotation. This holds for any instrument, as is the SOI, mounted on the telescope tube. Our ADC consists of a pair of fused silica prisms with the same apec angle and a constant 180 degree orientation offset. The forward prism corrects for atmospheric dispersion by moving longitudinally in the beam and the fixed prism corrects for the image plane tilt. An image shift occurs depending on the prism separation that can be compensated for by the telescope pointing model.

At present the ADC is not used during SOI observations.

The SOI Software
Analysis Computer

The data acquisition and preliminary reduction computer is called **soaric1**. A number of different ways to logon to this machine exist, depending upon your preference. These methods are discussed below.

- **From Cerro Pachón:**
  - The mountain staff or your support scientist will show you the computer on which you can obtain and analyze your data. To open the SOI user control panels:
    - Double click on the "Scroll Lock" Key
    - A window will open displaying the names of the computers to which you can connect
    - Use the "Up" and "Down" arrow keys to highlight "SOI 1 ic1" and press "ENTER/RETURN"
    - If the Data Acquisition and Analysis windows are running, you can skip to the [GUI Layout](#) section of this manual. If these GUIs are not running, skip to the [Starting and Stopping the Data Acquisition GUI](#) and [Starting and Stopping the Data Analysis GUI](#) sections of this Manual.

- **From the Remote Observing Center in La Serena:**
  - Log on to the SOI account using the username and password provided to you by the instrument scientist (Sean Points). If forgotten, these are posted on a list near the door.
  - Start the Data Acquisition GUI by typing the following command from a terminal command line on the GNU/Linux computer in the remote observing center:
    ```
    > vncviewer -Shared soaric1.ctio.noao.edu:9 &
    ```
    Log on to the vncviewer with the password provided by the instrument scientist.
  
  - Start the Data Analysis GUI by typing the following command from a terminal command line on the GNU/Linux computer in the remote observing center:
    ```
    > vncviewer -Shared soaric1.ctio.noao.edu:8 &
    ```
    Log on to the vncviewer with the password provided by the instrument scientist.

In most cases the GUIs should be started and you will be presented with a data acquisition screen and data analysis screen as shown in Figure 5.
Starting and Stopping the Data Acquisition GUI

If the data acquisition GUI has not been started, then one should see a light blue screen in the soaric1:9 VNC window with an xload and xclock window that has 8 buttons labeled FileMgr, terminal, SOI, IRAF, Gimp, gview, www, and some window manager utilities. To start the data acquisition software:

- Single click on the 'SOI' button. This will pop up the SOIGUI_GuiLogic window.
- On the SOIGUI_GuiLogic window, there are 3 buttons in the upper right-hand corner labeled Init Ctrl, Rel Ctrl, and EXIT.
- To the immediate left of these buttons are various instrument status indicators. If the startup has been successful, all of the instrument status lights should be a bright green. If any of these lights are red, the most probable cause is that the Leach Controller is not initialized. To initialize the controller, press the "Init Ctrl" button. After a few seconds, all of the indicator lights should be bright green.
- You are now ready to operate SOI.

If the data acquisition GUI needs to be stopped:

- Single click on the 'EXIT' button in the upper right-hand side of the data acquisition GUI.
- Under most normal cases, the GUI will close automatically and nothing else remains to be done.
- In certain cases, a window will appear and ask if you want to shutdown the LabView application. This window consists of a toggle switch labeled "Exit Labview". Click on the toggle switch so that it is in the down position. Then click on the "File" tab in that window and "Quit" LabView. After this, all of the LabView windows should close.

Starting and Stopping the Data Analysis GUI

The SOI data analysis VNC window (soaric1:8) has the same base layout as the SOI data acquisition VNC window (soaric1:9). If the data analysis windows are not up, you will see the xload window and the accompanying buttons described in the data acquisition GUI startup section. Single click on the IRAF button and an IRAF window will open. You will want to open a DS9 window from within IRAF and load the mscred IRAF package.

Basic GUI Layout

All observing with the SOAR Optical Imager (SOI) is handled through the Data Acquisition GUI. Upon successful startup of the SOI data acquisition GUI on soaric1:9, one should see that the SOI data acquisition window looks something like that shown in Figure 5.

The SOI observing GUI can be divided into certain distinct regions as shown in Figure 6. These include the:
Figure 6: The SOI data acquisition GUI with regions demarcated and labeled.

- **Connection Status Region** - This region of the GUI shows the local time in La Serena and the connection status of the instrument. After a successful startup of the GUI, all of the indicator lights toward the right-hand side should be a bright green. If they are not, you will first want to check if the Leach Controller has been initialized. You can initialize the Leach controller by clicking the "Init Ctrl" button. If lights remain red after initializing the controller, you may want to exit the GUI and restart. The Telescope Operations staff or your Observing Support Astronomer can help you with this.
• **Instrument and TCS Status Region** - This region shows various telemetry data from the instrument, as well as information obtained from the Telescope Control System (TCS). These data include the current RA and Dec, the airmass, the sidereal time, etc.

• **Exposure Status Region** - The Exposure Status Region provides information about the current exposure or sequence of exposures. If a sequence of exposures is being acquired, it shows how many of the current exposures have been completed in the uppermost status bar. Below that, there are counters that show how long the observation has been integrating and if the observation has been paused for any reason. Further down, this section of the GUI shows what the exposure time, i.e., how long the shutter will be open. Finally, this section of the GUI shows the readout and writeout progress of the current exposure.

• **CCD Geometry Region** - This region lies to the immediate right of the Exposure Status Region. Its primary purpose is to allow the observer to see the current readout mode and binning of the SOI CCDs. Furthermore, it alerts the observer if the a region-of-interest has been defined (i.e., only a section of the CCDs will readout), if the automatic image display feature is enabled for the SOI Data Analysis GUI, and if the observer is using the observing grid mode. If any of these features are enabled, the indicator lights in this section of the GUI will be bright green instead of the nominal dark green. These features will be discussed in more detail in the **Advanced Tools** section below.

• **Directory and Filename Status Region** - This region of the GUI lies below the Exposure Status and Geometry Status sections. They basically allow the observer to set the (1) directory in which the data will be written, (2) change the base filename of an image, and (3) change the sequence number of the observation. In standard practice, the directory path for data obtained by visiting NOAO observers is:

• **Data Acquisition Region** - This section of the Data Acquisition GUI lies toward the bottom of the display and essentially reflects the nuts and bolts of the image acquisition procedure. This section of the GUI allows the observer to select the **OBSTYPE** of an image. The currently defined OBSTYPEs are:
  - ZERO
  - OBJECT
  - DARK
  - DFLAT
  - SFLAT

These OBSTYPEs can be chosen by clicking on the appropriate tab in the upper left-hand side of this section of the GUI. In general, these tabs allow an observer to input the image title and an additional comment to the FITS header. All of the tabs that permit an observer to change the OBSTYPE of an image allow the observer to also change the exposure time of the image, except for a ZERO or Bias Frame observation. Furthermore, if the OBSTYPE is selected to be a DFLAT (Dome Flat), the observer is allowed to turn on and set the intensity of the White Spot Lamps. The Telescope Operators can provide you with the correct intensity values and exposure times for your Dome Flat observations. The "Misc" and "Focus" tabs in the Data Acquisition Region are not currently supported.

Also included in the Data Acquisition section of the SOI GUI are the Start, Pause, Stop, and Abort buttons. Pressing the "Start" button will begin the process of taking an exposure of the currently selected OBSTYPE. During an exposure, one may also click on the "Pause", "Stop", or "Abort"
buttons. The "Pause" button will suspend an ongoing exposure until the "Resume" button is clicked. During this time, the shutter should close so that the detectors are no longer integrating on the sky, but the detectors will continue to accumulate cosmic ray hits and dark current. Upon clicking the "Resume" button, the shutter will open and the exposure will proceed as scheduled. Clicking the "Stop" button during an exposure will close the shutter, readout the CCDs, and write the data to the disk. Clicking the "Abort" button will close the shutter and discard the current image. The "Abort" button does not write data to the disk.

Finally, on the right-hand side of the Data Acquisition Region, the observer can select the filter to used during the current observation. To change a filter:

- Click on the dialog box below the "Filters List"
- A list should open that will allow you to choose the filter to be used. Simply click on the filter you wish to use. A checkmark should appear on the left-hand side of the filter designation. As shown in Figure 7, the u-SDSS filter has been selected.
- After selecting the appropriate filter, click the filter "Move" button
- One should see the "Filter Status" display go to "Active" while the filter wheels are moving. Once the filter change is complete the "Current Filter" value should be the same as the "Filters List" value; the "Official Name" value should also reflect this change; and the "Filter Status" should change from "Active" to "Done".

- **Advanced Tools Region** - This section allows the observer more enhanced control of the SOAR Optical Imager. In this region the observer is able to use the:
  - The *Obs Editor* allows the observer to change the Proposal Specific data such as, Proposal ID,
Observation ID, Date, Institution, Proposal PI, and Observer, as shown in Figure 8. If you need to edit the values in the given fields, click on the "Edit" button. Once you have finished editing the fields, click on the "Apply" button and "Close" the window.

Figure 8: The SOI Observer Information interface.

- The Filter Editor changes the designations of the installed filters. At present, the TelOps staff will change these for you. Visiting observers SHOULD NOT use the Filter Editor. The filter editor is shown in Figure 9.
The Geometry Editor changes the readout section of the CCDs. It allows an observer to change the binning and select regions-of-interest. The standard SOI binning is 2x2 so that one has ~0.15"x0.15" pixels. The SOI Geometry Editor, shown in Figure 10, allows one to change binning and Region-of-Interest readout of the CCDs.
To change the binning, click on the Binning "Select" Box. The available options will be shown to you. To finalize your binning selection, click the "Apply" button to the right of the selection box. You should notice that Binning indicator in the Geometry Status section of the GUI flashes from yellow to red during this change. The background of the Binning indicator will remain white after the binning has successfully changed.

To change the Region-of-Interest (ROI) readout of the CCDs, one must first enable the ROI dialog by clicking on the "ROI Enable" button on the Geometry Editor screen. After pressing the ROI button, you can enter the detector coordinates (x,y) and the size in (x,y) of the ROI. The detector coordinates are related to the current binning. Thus, for a 2x2 binned observation, the center of the detector is at (x,y) = (1024,1024).

- The Grid Tool allows the Visiting Astronomer to use certain pre-defined dithering patterns either in pixel or RA/Dec coordinates. To use the Grid Tool, click on the Grid Tool button in the Advanced Tools Region of the GUI. Select the mode with which you want to execute your dither patterns (i.e., pixels or arcsec; Note: only RA and Dec offsets in arcsecs are supported at this time) by highlighting the appropriate tab in the upper left-hand corner of the Grid Tool, as seen the left-hand side of Figure 11. Choose a dither pattern by selecting the "Pattern" dialog box as shown in the right-hand side of Figure 11. The current pre-defined patterns are:
  - Pattern 0 (NW) - This dither pattern performs telescope offsets to the North and West of the current position.
  - Pattern 1 (SE) - This dither pattern performs telescope offsets to the South and East of the current position.
  - Pattern 2 (SW) - This dither pattern performs telescope offsets to the South and West of the current position.
  - Pattern 3 (NE) - This dither pattern performs telescope offsets to the North and East of the current position.
  - Pattern 4 (W) - This dither pattern performs telescope offsets to the West of the current position.
  - Pattern 5 (E) - This dither pattern performs telescope offsets to the East of the current position.
  - Pattern 6 (N) - This dither pattern performs telescope offsets to the North of the current position.
  - Pattern 7 (S) - This dither pattern performs telescope offsets to the South of the current position.
  - Pattern 8 (NSEW) - This dither pattern performs telescope offsets in a classic grid.
In order to apply a dither pattern to the Grid Tool, one should

- Select the Grid Tool mode **(arcsec)**.
- Select the grid pattern (see above).
- Enter the step size for the units selected above (e.g., 15" according to the left-hand side of Figure 11.).
- Enter the number of steps to complete. As this tool is currently implemented, the current position of the telescope is assumed to be the 0th step of the grid. That is, if as in the example shown in left-hand side of Figure 11, entering "2" into the "Step" box of the grid tool will perform a 3-point dither with the current telescope position being the 0th dither point, a 15" W offset being the 1st dither point, and another 15" W offset being the 2nd dither point. Thus, one thing to note is that subsequent dithers with this tool are relative and not absolute.
- In order to calculate the currently selected dither pattern, click on the "Enter" button below the grid pattern editor.
- To **Enable** the Grid Tool for observations, click on the "Enable" button in the bottom center of the Grid Tool.
- To **Apply** the currently selected dither pattern to the next observation, click the "Apply" button.

- That is to say, in order to perform a pre-defined dither pattern, one should: (1) select the arcsec Grid Tool mode; (2) determine the Grid Tool step size; (3) choose a grid pattern; (4) choose a number of grid steps; (5) click the "Enter" button; (6) Enable the Grid Tool (the grid light should light next to the left-hand side of the "Grid" indicator in the "Geometry Status" section of the GUI); and apply the grid for the observations that will be taken upon clicking the "Start" button.

- The **Script Tool** is currently undergoing Engineering tests. More information about this tool will be released at a later date.
- The **Offset and Engineering Tools** should not to be used by Visiting Astronomers at this time until more engineering tests are made. As such, nothing will be mentioned of them here.
Observing with SOI

Before Your Run

Prior to your run at SOAR with the Optical Imager, you should have completed the Instrument Setup Form [60]. When using SOI, it is important to send in this form well ahead of time so the proper filters [6] can be installed before your run. In the instrument setup form you can also specify what binning you will use during your run. The default readout is for 2x2 binning with fast readout. Information on the gain settings and readout noise for various binning options under fast readout can be found in the SOI Overview.

In addition to filling out the Instrument Setup Form [60], visiting observers should read the SOAR Visiting Astronomer’s webpage [61] for general information about traveling to and within Chile. Furthermore, visiting investigators should fill out the Travel Information Questionaire [62] so that your transportation and lodgings can be arranged.

Observing logs for your run can be downloaded here [63].

Setting Up for the Start of Your Night/Run

Before you begin observing with SOI, you should first make sure that the data acquisition GUI and the data analysis GUI are running as shown above in Figure 5. If these GUIs are not running, please refer to the sections of this manual about Starting and Stopping the Data Acquisition GUI [29] and Starting and Stopping the Data Analysis GUI [30]. If you have problems with starting either of these, please contact the Telescope Operators or the SOI instrument scientist (Sean Points). They will be able to help you with this task.

After the GUIs are running, you should check that:

- all of the connection status indicator lights are green
- the binning is set to the proper value (i.e., 1x1, 2x2, or 4x4). Information on how to set the binning is given in the Geometry Editor section if the Basic GUI Layout [31].
- the correct data acquisition directory is set. This directory must first be created on the file system before it can be set in the GUI. Refer to the Directory and Filename Status region of the Basic GUI Layout [31] section of this manual for more instructions on how to do this if you have problems.
- the filters you requested are installed. You can check which filters are installed by clicking on the Filters List window in the Data Acquisition section of the observing GUI. If the filters you requested are not installed, please contact the Telescope Operations Staff so that they may install the necessary filters.
- the proper observer data has been entered into the Obs Editor section of the data acquisition GUI. This can be done by clicking on the Obs Editor button in the Advanced Tools section of the GUI. After that information has been updated, click on the "Apply" button so that changes will go into effect.

You are now ready to use the SOAR Optical Imager. Information on how to takes exposures can be found in Basic GUI Layout [31] section.

You can also look at the step-by-step SOI Observer’s Cookbook [5] for help on observing with SOI.
After Your Night/Run

At the end of your observing night, please fill out the End-of-Night report for the telescope. Please make note of any problems that were encountered during the night so that they may be resolved before the next night's observing.

Also, at the end of your night observing with SOI, you may want to transfer your data back to your home institution. To do so, open a Terminal window in either the Data Acquisition GUI or the Data Analysis GUI. Once the Terminal window is open, change directories to where your data are located on soaric1. You may then use scp at the shell prompt to copy the data to your home institution.

After your run is complete, please fill out the End-of-Run report.

Evaluating and Reducing SOI Images

In this section we discuss the software and observing procedures needed for the following:

1. How to evaluate the observations as they are obtained at the telescope, including how to display SOI images and how to examine the delivered image quality,
2. Calibration data that should be obtained at the telescope, and
3. How to reduce the images.

Observers familiar with CCD cameras, images obtained using the NOAO Mosaic cameras, and the IRAF reduction and analysis software for the most part will find the processing of SOI images to be familiar. At the same time, there are some differences that we touch upon briefly here. To start with, SOI images are recorded in a special multi-extension FITS format (MEF). In brief, the SOI CCDs are saved as individual images grouped together as separate entities in a larger FITS file; only at the end of the reduction are the CCDs assembled as a single large astronomical image. Because of this special format, most IRAF routines will not work directly on the full SOI files. To provide for processing of the SOI format, as well as reduction and analysis tasks specific to SOI, we use some of the MSCRED IRAF routines that were developed for the Mosaic cameras in addition to having developed some SOI specific IRAF routines. Almost all of the software tasks that we discuss below presume that you will be working within this environment.

A key factor that drives both the data taking and reduction of SOI images is the presumption that the final astronomical exposure will be built from a number of SOI images obtained by dithering the telescope. This places strong demands on the quality of the data reduction to ensure the uniformity of the photometric response of the reduced image. The scripts we use for reducing SOI data can be found via FTP at SOAR.

If any help is needed to install these scripts or reduce your SOI data, please contact the SOI instrument scientist.

NOTE: These packages have been quasi-independently developed by astronomers at SOAR, CTIO and MSU. They contain slightly different versions of the same tasks. So it is recommended to download one or
the other packages, but not both. For non-expert IRAF MSCRED users, it is advised that the MSU routines be downloaded.

**Working with SOI/MEF Data Files**

An excellent summary of the MSCRED reduction routines is provided in the two guides written by Frank Valdes: [Mosaic Data Reduction System](#) [67] and the Guide to the NOAO Mosaic Data Handling System [68]. The latter link is available in the IRAF/MCSRED package by the command "help mscguide". We encourage SOI users to read through these documents before attempting to reduce their data for the first time. These guides also provide a thorough description of all MSCRED tasks that may be of use during the night's observing.

The SOAR Optical Imager data format is a multi-extension FITS (MEF) file. The file contains five FITS header and data units (HDU). The first HDU, called the primary or global header unit, contains only header information which is common to all the CCD images. The remaining four HDUs, called extensions, contain the header and images from the four amplifiers on the two CCDs.

The fact that the image data is stored as FITS format images is not particularly significant. A single FITS format image file may be treated in the same way as any other IRAF image format. The significant feature is the multi-extension nature of the data format. This means that commands that operate on images need to have the image or images within the file specified. Only commands specifically intended to operate on MEF files, such as those in the MSCRED package, can be used by simply specifying the file name. Commands that operate on files rather than images, such as copying a file, may be used on MEF files. In general, it is safest to use only MSCRED commands on MEF files. IRAF V2.11 is required to run MSCRED. The basic syntax for specifying an image in a MEF file to an IRAF task is:

```
filename.fits[extension]
```

where "filename" is the name of the file. The ".fits" extension does not need to be used. The "extension" is the name of the image. For the SOI data the 4 CCD images have the names "im1" through "im4" (but the simple "1" through "4" works, too). The extension position in the file (where the first extension is 1) may also be used. To access the global header (for listing or editing) the extension number is 0; i.e. filename.fits[0].

There is currently no wildcard notation for specifying a set of extensions. So to apply an arbitrary IRAF command that takes a list of images you must either prepare an @list (or type the list explicitly) or use the special MSCCMD command. The task MSCCMD takes an IRAF command with the image list parameter replaced by the special string "$input". The input list of SOI files will then be expanded to a list of image extensions.

**Displaying and Evaluating SOI Images at the Telescope**

During observing, a small set of IRAF commands are commonly used to examine the data. This section describes these common commands. While this section is oriented to examining the data at the telescope during the course of observing, the tools described here are also used when reducing and analyzing the data at a later time.
The two IRAF commands DISPLAY and MSCDISPLAY are used to display the images in DS9. DS9 and an IRAF command window should be running on the data analysis GUI [30]. The DISPLAY task is used to display individual images - in this context, the individual amplifiers in a SOI exposure designated by the appropriate extension ID. There are many display options that are discussed in the help page that will not be discussed here. The only special factor in using this task with the SOI data is that you must specify which image to display using the image extension syntax discussed previously. As an example (see Figure 12), the central portion of extension im2 (i.e., AMP#2) in displayed in the first frame of a DS9 window using the IRAF DISPLAY command and the whole image is displayed in the second frame of a DS9 window using the IRAF MSCDISPLAY command:

- mscred> disp obj.064.fits[2] 1 zs+
- mscred> mscdisp obj.064.fits 2 zs+

![Figure 12: (left) A DS9 image display of the second SOI amplifier using the IRAF DISPLAY command. (right) A DS9 image display of an entire image using the IRAF MSCDISPLAY command.](image)

The MSCDISPLAY task is based on DISPLAY with a number of specialized enhancements for displaying MEF data. It displays the entire MEF observation in a single frame by "filling" each image in a tiled region of the frame buffer. The default filling (defined by the order parameter) subsamples the image by uniform integer steps to fit the tile and then replicates pixels to scale to the full tile size. The resolution is set by the frame buffer size defined by the "stdimage" variable.

Many of the parameters in MSCDISPLAY are the same as DISPLAY and there are also a few that are specific to the task of displaying a mosaic of CCD images. The mapping of the pixel values to gray levels includes the same automatic or range scaling algorithms as in DISPLAY. This is done for each image in the mosaic separately. The new parameter "zcombine" then selects whether to display each image with its own display range ("none") or to combine the display ranges into a single display range based on the minimum and maximum values ("minmax"), the average of the minimum and maximum values.
"average"), or the median of the minimum and maximum values. The independent scaling may be most appropriate for raw data while the "minmax" scaling is recommend for processed data. Another new optional answer here is "auto", which is the default, will try to use the best option, given the status of the data.

During your nightly observing, MSCDISPLAY is normally used during image acquisition to insure that your object is centered at the desired location and that it does not lie in the chip gap between the CCDs. In order to measure the delivered image quality of an exposure, one normally uses the IRAF DISPLAY and IMEXAM commands to measure the FWHM and the ellipticity of stars in the field. The FWHM can be compared to the seeing reported by the Cerro Pachón seeing monitor [69]. If the measured image quality is considerably worse than the site seeing, you may want to re-tune the mirror. Another measure of the image quality is the ellipticity of stars in the observed field. If the ellipticity, as measured by the IMEXAM command, is greater than 0.15 you may want to re-tune the mirror. The telescope operators normally tune the mirror at the beginning of the night. After that, it is the observer’s responsibility to check the image quality and request that the mirror be re-tuned.

**Calibration Data to Obtain at the Telescope**

**In general the only calibration data that one needs to obtain at the telescope are Bias frames, and Flats (Dome and/or Twilight Sky).** These data are obtained by clicking on the appropriate tab in the Data Acquisition Region of the [Data Acquisition GUI][31]. This will set the proper "OBSTYPE" keyword (i.e., ZERO, DOMEFLAT, or SKYFLAT) in the FITS headers for your calibration data.

[Click here for table of exposure times and lamp intensities][7] that will allow you to take dome flats with a peak intensity of ~20000 counts for most filters. When taking twilight sky flats, one should start observing as soon as you can after sunset, and take at least 5 twilight sky flats per filter. As soon as the peak counts for a particular filter are below ~30000, we recommend taking a series of exposures until either the peak counts drop below ~10000 or you have enough twilight sky flats to reduce your data.

**Fringing with SOI**

The only other calibration data one will likely need to obtain at the telescope are sky frames for fringing correction. As any thinned, back illuminated chip the SOI E2V devices suffer from considerable fringing in the Bessel I, SDSS-i and SDSS-z bands [70]. Fringing is most evident in longer exposures. Depending on the nature of your science targets you can take either of two approaches to approach obtaining the fringing correction images:

1) If your field is sparse and contains mostly/only stars (maybe a few faint, small galaxies), then you can use the same science frames to construct the master fringe removal frame. All that is needed is to carry out the observation as a series of dithered/offset observations, which you will later median-combine without registering (after the basic bias subtraction and flatfielding). This will eliminate all stars and leave only the sky background with the fringe pattern - see the examples in this page. [71] Just make sure you dither/offset by at least 10 arcsec between each exposure.

2) If you are imaging an extended object which covers a significant part of your field, like a large galaxy, or you are looking at a very crowded stellar field, the previous approach will not work. In this case you should offset to a nearby "empty" field (sparse stellar field) to take your sky data. Again, you need to split your total exposure time among at least 3, preferably 5 exposures, that you will then median-
combine; dither at least 10 arcsec between each frame. Ideally, your total "empty" sky image should have an integration time equal to your science frame. It is important to bear in mind the fringing pattern will not be the same across all the sky or all position angles, so for the best fringe removal, you should create the master fringe frame using either the same science data or images obtained in the same general area of the sky. Removal of the fringing pattern is done by applying the master fringe frame as an illumination correction to the data that have been already bias-subtracted and flatfielded. More details can be found in these links on fringing correction:

- Basis for a SOAR Imager pipeline (Maia et al 2011) [72]
- https://www.eso.org/sci/facilities/lasilla/instruments/efosc/inst/fringing.html [73]
- http://stsdas.stsci.edu/cgi-bin/gethelp.cgi?rmfringe [75]

Basic Data Reduction

This section assumes that either the NOAO SOI reduction package or the MSU SOI reduction package has been downloaded and successfully installed on your local reduction computer. If not, please contact the SOI instrument scientist [11] (Sean Points) for help.

The first step in reducing SOI data should be to perform the BIAS scan and TRIM section corrections. The values for BIASSEC and TRIMSEC are given in the FITS headers for each amplifier. Please see the IRAF MSCRED CCDPROC help pages for further information. After all of the images are Bias-subtracted and Trimmed, the Zero (or Bias) frames should be combined into a master bias for the night using the IRAF mscred.zerocombine task. After this step, the flats and object images should be Zero-corrected by setting the appropriate flags in mscred.ccdproc. Combine the dome flats and if necessary twilight flats, depending on the filter. Apply the Flat-field correction for your data using mscred.ccdproc. You should now be ready to perform photometry on your SOI data.

Step-by-step instructions on reducing SOI data can be found here [8].

Dome Flat - Typical Exposure Times

This page contains the Typical Exposured Times used in Dome Flats for SOI and SAMI. The values are not precise but are meant to give the astronomer an idea.

Click here for a table. [7]
SOI Image Reduction

SOI Notes

S. Points, 12 Oct 2011.

In the following document GNU/Linux and IRAF commands are preceded by "linux-prompt>" and "IRAFPackage>", respectively.
You may want to download soar.cl [76], soinoao.tar [77], and soimsu.tar.gz [78] first.

1) Start IRAF on your local machine

```
linux-prompt> cl - ecl
ecl> soar
soar> soinoao
soinoao> mscred
mscred>
```

Get filter information from the FITS header. Please remember that SOI can hold up to eight (8) filters in two (2) filter wheels at any given time. This can be done with IRAF using the TABLES keypar package, but can be cumbersome. An example of an IRAF script that can do this can be found at get_soi_filter_keyword.cl [79].
This task is more easily done by using the WCSTools [80] gethead program. If the WCSTools [80] are installed on your local machine, you can determine the SOI filter information by entering the following on a command line

```
linux-prompt> gethead FILPOS FILTERS FILTER1 FILTER2 *fits
```

This returns the following information:

```
dflat.001.fits   1 5 s0000_s0002 s0000 Open s0002 B Bessell
... 
dflat.011.fits   1 4 s0000_s0003 s0000 Open s0003 V Bessell
... 
dflat.021.fits   1 3 s0000_s0004 s0000 Open s0004 R Bessell
... 
dflat.031.fits   1 2 s0000_s0005 s0000 Open s0005 I Bessell
... 
sflat.001.fits   2 1 s0001_s0000 s0001 U Bessell s0000 Open
```

The filter position is used to record the filter name in the images headers and is required to handle image subsets. After the filter information is determined.

From the above example with "gethead", we can see that:
U_Bessell is in filpos 2 1
B_Bessell is in filpos 1 5
V_Bessell is in filpos 1 4
R_Bessell is in filpos 1 3
I_Bessell is in filpos 1 2

Note: Position 1 in each filter wheel is normally designated as "Open".

2) Use IRAF to create the necessary subsets

mscred> hsel *.fits[1] $I 'filpos == "2 1"' > listU
mscred> hsel *.fits[1] $I 'filpos == "1 5"' > listB
mscred> hsel *.fits[1] $I 'filpos == "1 4"' > listV
mscred> hsel *.fits[1] $I 'filpos == "1 3"' > listR
mscred> hsel *.fits[1] $I 'filpos == "1 2"' > listI

Trim the last eight (8) characters of the filenames, i.e., ".fits[1]". This can be done easily using the following GNU/Linux commands:

linux-prompt> cp listU listU2
linux-prompt> vi listU2

After you can edit the list file, perform a global search and replace of the .fits[1] characters using the following:

:1,$ s/.fits\[1\]///g

Repeat this for other filters.

If you are not familiar with vi, you can trim the last 8 characters of the filename using the following IRAF code on the command line:

mscred> int size_str
list="listU"
while (fscan (list,s1) ! = EOF) {
    size_str=strlen(s1)
    print substr(s1,1,size_str-8) >> listU2
}

Repeat this for other filters.

Update the "FILTER" parameter in the image headers within IRAF

mscred> soiupfilter @listU2 filter_=yes filter=U
mscred> soiupfilter @listB2 filter_=yes filter=B
mscred> soiupfilter @listV2 filter_=yes filter=V
mscred> soiupfilter @listR2 filter_=yes filter=R
mscred> soiupfilter @listI2 filter_=yes filter=I

Check the IMTYPE keyword and the Filter subsets

mscred> ccdlist *fits[1]
This should give something like the following:

dflat.001.fits[1][563,2048][ushort][FLAT][1][B]:Dome flat
... 
dflat.011.fits[1][563,2048][ushort][FLAT][1][V]:Dome flat
... 
dflat.021.fits[1][563,2048][ushort][FLAT][1][R]:Dome flat
... 
dflat.031.fits[1][563,2048][ushort][FLAT][1][I]:Dome flat
... 
sflat.001.fits[1][563,2048][ushort][SKYFLAT][1][U]:sflat
... 
obj.006.fits[1][563,2048][ushort][OBJECT][1][U]:T Phe
... 
obj.007.fits[1][563,2048][ushort][OBJECT][1][B]:T Phe
... 
obj.008.fits[1][563,2048][ushort][OBJECT][1][V]:T Phe
... 
obj.009.fits[1][563,2048][ushort][OBJECT][1][I]:T Phe

3) Check the BIASSEC, DATASEC, and TRIMSEC parameters of the FITS files.
We have received reports that these may be off by a pixel. Thus, it is always a good idea to inspect the
data by hand before moving on. The latest and greatest values for these parameters that we have are...

For Pre 14 Nov 2008 Data (SOI Leach II Controller, 2 2kx4k CCDs):

<table>
<thead>
<tr>
<th>AMP</th>
<th>BIASSEC</th>
<th>TRIMSEC</th>
<th>DATASEC</th>
</tr>
</thead>
<tbody>
<tr>
<td>1</td>
<td>[541:562,1:2048]</td>
<td>[29:540,1:2048]</td>
<td>[29:540,1:2048]</td>
</tr>
<tr>
<td>3</td>
<td>[541:562,1:2048]</td>
<td>[29:540,1:2048]</td>
<td>[29:540,1:2048]</td>
</tr>
</tbody>
</table>

For 14 Nov 2008 -- DD MMM 2010 Data (SOI Leach III Controller, 2 2kx4k CCDs):

<table>
<thead>
<tr>
<th>AMP</th>
<th>BIASSEC</th>
<th>TRIMSEC</th>
<th>DATASEC</th>
</tr>
</thead>
<tbody>
<tr>
<td>1</td>
<td>[541:568,1:2048]</td>
<td>[29:540,1:2048]</td>
<td>[29:540,1:2048]</td>
</tr>
<tr>
<td>2</td>
<td>[1:28,1:2048]</td>
<td>[29:540,1:2048]</td>
<td>[29:540,1:2048]</td>
</tr>
<tr>
<td>3</td>
<td>[541:568,1:2048]</td>
<td>[29:540,1:2048]</td>
<td>[29:540,1:2048]</td>
</tr>
<tr>
<td>4</td>
<td>[1:28,1:2048]</td>
<td>[29:540,1:2048]</td>
<td>[29:540,1:2048]</td>
</tr>
</tbody>
</table>

For Post DD MMM 2010 Data (SOI Leach III Controller, 1 4kx4k CCD):

<table>
<thead>
<tr>
<th>AMP</th>
<th>BIASSEC</th>
<th>TRIMSEC</th>
<th>DATASEC</th>
</tr>
</thead>
<tbody>
<tr>
<td>1</td>
<td>TBD</td>
<td>TBD</td>
<td>TBD</td>
</tr>
<tr>
<td>2</td>
<td>TBD</td>
<td>TBD</td>
<td>TBD</td>
</tr>
<tr>
<td>3</td>
<td>TBD</td>
<td>TBD</td>
<td>TBD</td>
</tr>
<tr>
<td>4</td>
<td>TBD</td>
<td>TBD</td>
<td>TBD</td>
</tr>
</tbody>
</table>

If the BIASSEC, TRIMSEC, and DATASEC keywords are not correct, you will need to change them. It is
easy to do this within IRAF using the mscred.ccdhedit task. Be sure to use the appropriate values
depending upon the date of your observations.
4) Combine the bias frames

mscred> epar zerocombine

I R A F

Image Reduction and Analysis Facility

PACKAGE = mscred
TASK = zerocombine

input    = zero* List of zero level images to combine
(output = Zero) Output zero level name
(combine= median) Type of combine operation
(reject = avsigclip) Type of rejection
(ccdtype= ZERO) CCD image type to combine
(process= yes) Process images before combining?
(delete = no) Delete input images after combining?
(scale = none) Image scaling
(statsec= ) Image section for computing statistics
(nlow = 0) minmax: Number of low pixels to reject
(nhigh = 1) minmax: Number of high pixels to reject
(nkeep = 1) Minimum to keep (pos) or maximum to reject (neg)
(mclip = yes) Use median in sigma clipping algorithms?
(isigma = 3.) Lower sigma clipping factor
(hsigma = 3.) Upper sigma clipping factor
(rdnoise= 4.4) ccdclip: CCD readout noise (electrons)
(gain = 2) ccdclip: CCD gain (electrons/DN)
(snoise = 4.4) ccdclip: Sensitivity noise (fraction)
(pclip = -0.5) pclip: Percentile clipping parameter
(blank = 0.) Value if there are no pixels
(mode = q)

mscred> epar ccdproc

I R A F

Image Reduction and Analysis Facility

PACKAGE = mscred
TASK = ccdproc

images  = zero*fits List of Mosaic CCD images to process
(output = ) List of output processed images
(bpmaps= ) List of output bad pixel masks
(ccdtype= ZERO) CCD image type to process
(noproc = no) List processing steps only?
(xtalkco= no) Apply crosstalk correction?
(fixpix = no) Apply bad pixel mask correction?
(oversca= yes) Apply overscan strip correction?
(trim = yes) Trim the image?
(zero-cor= no) Apply zero level correction?
(darkcor= no) Apply dark count correction?
(flatcor= no) Apply flat field correction?
(sflatco= no) Apply sky flat field correction?
(split = no) Use split images during processing?
(merge = no) Merge amplifiers from same CCD?
(xtalkfi= ) Crosstalk file
(fixfile= ) List of bad pixel masks
(saturat= 45000) Saturated pixel threshold
(sgrow = 0) Saturated pixel grow radius
(bleed = INDEF) Bleed pixel threshold
(btrail = 10) Bleed trail minimum length
(bgrow = 0) Bleed pixel grow radius
(biassec= image) Overscan strip image section
(trimsec= image) Trim data section
(zero = Zero.fits) List of zero level calibration images
(dark = ) List of dark count calibration images
(flat = ) List of flat field images
(sflat = ) List of secondary flat field images
(minrepl= 1.) Minimum flat field value
(interac= no) Fit overscan interactively?
(funcio= legendre) Fitting function
(order = 2) Number of polynomial terms or spline pieces
(sample = *) Sample points to fit
(naverag= 1) Number of sample points to combine
(niterat= 4) Number of rejection iterations
(low_rej= 3.) Low sigma rejection factor
(high_rej= 3.) High sigma rejection factor
(grow = 0.) Rejection growing radius
(fd = )
(fd2 = )
(mode = q)

After you have checked all of the parameters, you can run zerocombine. It is generally useful to fit the first few overscan regions interactively until you are satisfied with the parameters.

5) Create your master flat fields in each filter

mscred> epar flatcomb

I R A F
Image Reduction and Analysis Facility
PACKAGE = mscred
TASK = flatcombine

input = dflat*fits List of flat field images to combine
(output = Dflat_) Output flat field root name
(combine= median) Type of combine operation
(reject = avsigclip) Type of rejection
(ccdtype= ) CCD image type to combine
(process= yes) Process images before combining?
(subsets= yes) Combine images by subset parameter?
(delete = no) Delete input images after combining?
(scale = mode) Image scaling
(statsec= ) Image section for computing statistics
(nlow = 1) minmax: Number of low pixels to reject
(nhigh = 1) minmax: Number of high pixels to reject
(nhigh = 1) minmax: Number of high pixels to reject
(nkeep = 1) Minimum to keep (pos) or maximum to reject (neg)
(mclip = yes) Use median in sigma clipping algorithms?
(lsigma = 3.) Lower sigma clipping factor
(hsigma = 3.) Upper sigma clipping factor
(rdnoise= 4.4) ccdclip: CCD readout noise (electrons)
(gain = 2) ccdclip: CCD gain (electrons/DN)
(snoise = 0.) ccdclip: Sensitivity noise (fraction)
(pclip = -0.5) pclip: Percentile clipping parameter
(blank = 1.) Value if there are no pixels
(mode = q)

mscred> epar ccdproc

I R A F
Image Reduction and Analysis Facility
PACKAGE = mscred
TASK = ccdproc

images = dflat*fits List of Mosaic CCD images to process
(output = ) List of output processed images
(bpmasks= ) List of output bad pixel masks
(ccdtype= FLAT) CCD image type to process
(noproc = no) List processing steps only?
(xtalkco= no) Apply crosstalk correction?
(fixpix = no) Apply bad pixel mask correction?
(oversca= yes) Apply overscan strip correction?
(trim = yes) Trim the image?
(zero cor= yes) Apply zero level correction?
(darkcor= no) Apply dark count correction?
(flatcor= no) Apply flat field correction?
(sflatco= no) Apply sky flat field correction?
(split = no) Use split images during processing?
(merge = no) Merge amplifiers from same CCD?
(xtalkfi= ) Crosstalk file
(fixfile= ) List of bad pixel masks
(saturat= 45000) Saturated pixel threshold
(sgrow = 0) Saturated pixel grow radius
(bleed = INDEF) Bleed pixel threshold
(btrail = 10) Bleed trail minimum length
(bgrow = 0) Bleed pixel grow radius
(biassec= image) Overscan strip image section
(trimsec= image) Trim data section
(zero = Zero.fits) List of zero level calibration images
(dark = ) List of dark count calibration images  
(flats = ) List of flat field images  
(sflats = ) List of secondary flat field images  
(minrepl = 1.) Minimum flat field value  
(interact = no) Fit overscan interactively?  
(functio = legendre) Fitting function  
(order = 2) Number of polynomial terms or spline pieces  
(sample = 
  *) Sample points to fit  
(naverag = 1) Number of sample points to combine  
(niterat = 4) Number of rejection iterations  
(low_rej = 3.) Low sigma rejection factor  
(high_re = 3.) High sigma rejection factor  
(grow = 0.) Rejection growing radius  
(fd = )  
(fd2 = )  
(mode = q)  

6) Process your science frames for overscan, trim-correction, zero-subtraction, and flat field corrections.  
In general you will want to make a list of your data for each filter. Then you will want to apply a WCS to your data and create a mosaicked image. See the Mosaic Data reduction guide, by Frank Valdes [68].

This section is under construction.

SOI Technical pages

Technical information on SOI mostly of interest to SOAR technical and support staff.

- Distances from CCD to various optical train components [81]

Source URL: http://www.ctio.noao.edu/soar/content/soar-optical-imager-soi

Links